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Three-dimensional non-relativistic supergravity and torsion

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Abstract In this paper we present a torsional non-relativistic Chern–Simons (super)gravity theory in three spacetime dimensions. We start by developing the non-relativistic limit of the purely bosonic relativistic teleparallel Chern–Simons formulation of gravity. On-shell the latter yields a non-Riemannian setup with non-vanishing torsion, which, at non-relativistic level, translates into a non-vanishing spatial torsion sourced by the cosmological constant. Then we consider the three-dimensional relativistic $\mathcal{N} = 2$ teleparallel Chern–Simons supergravity theory and obtain its non-relativistic counterpart by exploiting a Lie algebra expansion method. The non-relativistic supergravity theory is characterized, on-shell, by a non-vanishing spatial super-torsion, again sourced by the cosmological constant.

1 Introduction

The introduction of torsion in a gravity theory can be performed through the teleparallel formulation of gravity [1–5], which is considered to be equivalent to General Relativity. Nevertheless, teleparallel gravity is characterized by a vanishing curvature and a non-vanishing torsion describing a non-Riemannian geometry denoted as Weizenböck geometry. In three spacetime dimensions, the role of torsion in black hole solutions and boundary dynamics has been of particular interest [6–13]. As it was shown in [6–8], torsional three-dimensional gravity theory possesses a BTZ-like black hole solution whose thermodynamic properties have been studied in [14–16]. Moreover, both teleparallel gravity and AdS gravity have identical asymptotic symmetries given by two copies of the Virasoro algebra [9]. Interestingly, three-dimensional teleparallel gravity can be formulated as a gauge

theory using the Chern–Simons (CS) formalism and considering the so-called teleparallel algebra [17]. The teleparallel CS gravity action can be alternatively recovered as a particular case of the Mielke–Baekler gravity model [18] by fixing the Mielke–Baekler parameters. Remarkably, the CS approach allows us to define a supersymmetric teleparallel gravity in the presence of $\mathcal{N} = p + q$ supercharges which in the vanishing cosmological constant limit reproduces the \mathcal{N} -extended Poincaré supergravity theory [17].

On the other hand, non-relativistic (NR) theories have received a growing interest due to their relation to condensed matter systems [19–30] and NR effective field theories [31–35]. A NR theory can be obtained by taking the limit $c \rightarrow \infty$, where c is the speed of light, of a relativistic theory. At the gravity level, the corresponding NR counterpart of the AdS spacetime is the Newton–Hooke symmetry which in the vanishing cosmological constant limit reproduces the Galilei symmetry [36–42]. In three spacetime dimensions, the CS formalism allows us to construct NR gravity actions whose underlying symmetry can be obtained as a NR limit of a relativistic algebra. However, it is necessary to consider additional $u(1)$ generators in order to avoid infinities and degeneracy in the NR limit. Indeed, the Poincaré algebra requires the presence of two extra $u(1)$ generators in order to apply an appropriate and well-defined NR limit with a non-degenerate bilinear form and reproduces the extended Bargmann algebra [43–47]. In the presence of a cosmological constant, the NR version of the $\text{AdS} \oplus u(1)^2$ symmetry corresponds to an extension of the Newton–Hooke symmetry [48–51].

The formulation of a NR supergravity theory has only been approached recently and remains as a challenging task mainly motivated by the diverse applications of these models in the context of holography and relativistic field theory [52–57]. One way to circumvent the difficulty to establish a well-defined NR limit in the presence of supersymmetry is through the Lie algebra expansion methods of [58–61], which allows us to obtain the respective NR version of

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where W^A is the spin connection and E^A denotes the vielbein satisfying $E_\mu^A E_\nu^B g^{\mu\nu} = \eta^{AB}$. The curvature two-form $F = dA + \frac{1}{2}[A, A]$ reads

$$F = R^A(W) J_A + R^A(E) P_A, \tag{2.4}$$

where

$$\begin{aligned} R^A(W) &= dW^A + \frac{1}{2}\epsilon^{ABC} W_B W_C, \\ R^A(E) &= T^A - \frac{1}{\ell}\epsilon^{ABC} E_B E_C. \end{aligned} \tag{2.5}$$

Here R^A is the Lorentz curvature two-form and T^A denotes the usual torsion two-form $T^A = dE^A + \epsilon^{ABC} W_B E_C$. The CS gravity action invariant under the teleparallel algebra (2.1) can be obtained considering the one-form gauge connection A (2.3) and the non-vanishing components of the invariant bilinear form (2.2) in the general expression for the CS action:

$$I_{CS}[A] = \frac{k}{4\pi} \int_{\mathcal{M}} \left\langle AdA + \frac{2}{3}A^3 \right\rangle. \tag{2.6}$$

Then, the teleparallel CS action reads

$$\begin{aligned} I_{TG} &= \frac{k}{4\pi} \int_{\mathcal{M}} \left\{ \alpha_0 \left(W^A dW_A + \frac{1}{3}\epsilon^{ABC} W_A W_B W_C \right) \right. \\ &\quad \left. + \alpha_1 \left(2E_A R^A(W) + \frac{4}{3\ell^2}\epsilon^{ABC} E_A E_B E_C - \frac{2}{\ell}T^A E_A \right) \right\}, \end{aligned} \tag{2.7}$$

where k is the CS coupling constant and is related to the gravitational constant G as $k = 1/(4\pi G)$. The teleparallel CS action is characterized by the Einstein–Hilbert term, the cosmological constant term plus a torsional term along the α_1 constant. On the other hand, the sector proportional to α_0 is known as the exotic CS term [89] and, as was mentioned in [17], it can be made to vanish without loss of generality. Nevertheless, for $\alpha_0 \neq 0$, the non-degeneracy of the invariant tensor (2.2) requires $\alpha_0 \neq -\frac{\ell}{2}\alpha_1$ and $\alpha_1 \neq 0$. One can then see that, for non-degenerate invariant tensor, the equation of motions are given by the vanishing of the curvature two-form (2.5). Unlike AdS gravity, which describes a Riemannian spacetime, the teleparallel CS gravity is characterized by a non-Riemannian geometry² in which the cosmological constant can be seen as a source for the torsion,

$$T^A - \frac{1}{\ell}\epsilon^{ABC} E_B E_C = 0. \tag{2.8}$$

In fact, the latter is an on-shell constraint for the theory. In the vanishing cosmological constant limit $\ell \rightarrow \infty$, the theory reduces to the usual Poincaré CS gravity theory with vanishing torsion.

² See, e.g., [90,91] for a recent review on non-Riemannian structures and applications in (super)gravity. Besides, we refer the reader to, e.g., [92–96] for further implications of a non-vanishing torsion and non-Riemannian structures in gravity and supergravity theories.

Let us note that the teleparallel gravity action can be recovered alternatively as a particular case of the Mielke–Baekler gravity model [18,97,98], whose action reads

$$I_{MB} = aI_1 + \Lambda I_2 + \beta_3 I_3 + \beta_4 I_4, \tag{2.9}$$

where a, Λ, β_3 and β_4 are constants and

$$\begin{aligned} I_1 &= 2 \int E_A R^A(W), \\ I_2 &= -\frac{1}{3} \int \epsilon_{ABC} E^A E^B E^C, \\ I_3 &= \int W^A dW_A + \frac{1}{3}\epsilon^{ABC} W_A W_B W_C, \\ I_4 &= \int E_A T^A. \end{aligned} \tag{2.10}$$

Indeed, by fixing the Mielke–Baekler parameters as follows,

$$\begin{aligned} a &= \frac{\alpha_1}{16\pi G}, \quad \Lambda = -\frac{\alpha_1}{4\pi G \ell^2}, \\ \beta_3 &= \frac{\alpha_0}{16\pi G}, \quad \beta_4 = -\frac{\alpha_1}{8\pi G \ell}, \end{aligned} \tag{2.11}$$

we recover the teleparallel CS gravity action (2.7) for $k = 1/(4\pi G)$.

3 Non-relativistic gravity with torsion

In this section we present the explicit construction of a non-relativistic version of the teleparallel CS gravity theory previously discussed. The NR action obtained is characterized by vanishing curvatures for the NR analogue of the spin connection and non-vanishing spatial torsion. Interestingly, the spatial component of the torsion is zero in the vanishing cosmological constant limit $\ell \rightarrow \infty$ and the NR gravity theory reproduces the known extended Bargmann gravity. The novel NR gravity theory presented here will be denoted as torsional NR gravity and offers us an alternative way to introduce a cosmological constant into the three-dimensional extended Bargmann gravity model.

3.1 $U(1)$ -enlargements

A torsional NR algebra can be obtained by applying the NR limit to a particular $U(1)$ -enlargement of the teleparallel algebra (2.1). To this end, let us introduce two extra $U(1)$ gauge fields to the field content. Then, the gauge connection one-form reads

$$A = W^A J_A + E^A P_A + M Y_1 + S Y_2. \tag{3.1}$$

As we shall see, the presence of these additional Abelian generators is required to ensure a finite and non-degenerate NR action in the limit process. Such extra gauge generators

yield the following non-vanishing components of the invariant tensor:

$$\langle Y_2 Y_2 \rangle = \alpha_0, \quad \langle Y_1 Y_2 \rangle = \alpha_1, \quad \langle Y_1 Y_1 \rangle = -\frac{2\alpha_1}{\ell}. \tag{3.2}$$

Then, the [teleparallel]⊕u(1)² algebra admits the non-vanishing components of the invariant tensor given by (2.2) along with (3.2). One can show that the relativistic CS action based on the [teleparallel]⊕u(1)² algebra is obtained considering the gauge connection one-form (3.1) and the invariant tensor (2.2) and (3.2) in the general expression of the CS action (2.6),

$$I_R = \frac{k}{4\pi} \int_{\mathcal{M}} \left\{ \alpha_0 \left(W^A dW_A + \frac{1}{3} \epsilon^{ABC} W_A W_B W_C + SdS \right) + \alpha_1 \left(2E_A R^A(W) + \frac{4}{3\ell^2} \epsilon^{ABC} E_A E_B E_C - \frac{2}{\ell} T^A E_A + 2MdS - \frac{2}{\ell} M dM \right) \right\}. \tag{3.3}$$

Let us note that we recover the relativistic CS action for the [Poincaré]⊕u(1)² symmetry in the flat limit $\ell \rightarrow \infty$.

3.2 Torsional non-relativistic gravity theory

The NR counterpart of the relativistic [teleparallel]⊕u(1)² algebra can be derived through an Inönö–Wigner procedure [99]. To this end, we shall first consider the decomposition of the *A* index as follows:

$$A \rightarrow (0, a), \quad a = 1, 2. \tag{3.4}$$

Then, we will introduce a dimensionless parameter ξ which will allow us to identify the relativistic generators in terms of the NR ones (denoted with a tilde) as:

$$J_0 = \frac{\tilde{J}}{2} + \xi^2 \tilde{S}, \quad J_a = \xi \tilde{G}_a, \quad Y_2 = \frac{\tilde{J}}{2} - \xi^2 \tilde{S}, \\ P_0 = \frac{\tilde{H}}{2\xi} + \xi \tilde{M}, \quad P_a = \tilde{P}_a, \quad Y_1 = \frac{\tilde{H}}{2\xi} - \xi \tilde{M}. \tag{3.5}$$

Furthermore, in order to ensure a finite NR symmetry after the contraction process we shall consider the following scaling of the length parameter:

$$\ell \rightarrow \xi \ell. \tag{3.6}$$

Then, considering the identification (3.5) and the limit $\xi \rightarrow \infty$ in the relativistic commutation relations (2.1), we obtain a novel NR symmetry spanned by the set of generators $\{\tilde{J}, \tilde{G}_a, \tilde{S}, \tilde{H}, \tilde{P}_a, \tilde{M}\}$ which satisfy the following commutators:

$$[\tilde{J}, \tilde{G}_a] = \epsilon_{ab} \tilde{G}_b, \quad [\tilde{G}_a, \tilde{G}_b] = -\epsilon_{ab} \tilde{S}, \quad [\tilde{H}, \tilde{G}_a] = \epsilon_{ab} \tilde{P}_b, \\ [\tilde{J}, \tilde{P}_a] = \epsilon_{ab} \tilde{P}_b, \quad [\tilde{G}_a, \tilde{P}_b] = -\epsilon_{ab} \tilde{M}, \quad [\tilde{H}, \tilde{P}_a] = -\frac{2}{\ell} \epsilon_{ab} \tilde{P}_b, \\ [\tilde{P}_a, \tilde{P}_b] = \frac{2}{\ell} \epsilon_{ab} \tilde{M}, \tag{3.7}$$

where we have defined $\epsilon^{ab} \equiv \epsilon^{0ab}$ and $\epsilon_{ab} \equiv \epsilon_{0ab}$. As in the extended Newton–Hooke symmetry [48,50], the torsional NR algebra (3.7) consists of spatial translations \tilde{P}_a , spatial rotations \tilde{J} , Galilean boosts \tilde{G}_a , time translations \tilde{H} and two central charges \tilde{S} and \tilde{M} . Let us note that the extended Bargmann algebra [54] is obtained in the vanishing cosmological constant limit $\ell \rightarrow \infty$. On the other hand, the NR symmetry (3.7) can be written as two copies of the Nappi–Witten algebra [100,101],

$$[\tilde{J}^\pm, \tilde{G}_a^\pm] = \epsilon_{ab} \tilde{G}_b^\pm, \\ [\tilde{G}_a^\pm, \tilde{G}_b^\pm] = -\epsilon_{ab} \tilde{S}^\pm, \tag{3.8}$$

by considering the following redefinitions:

$$\tilde{G}_a = \tilde{G}_a^+ + \tilde{G}_a^-, \quad \tilde{P}_a = -\frac{2}{\ell} \tilde{G}_a^-, \\ \tilde{J} = \tilde{J}^+ + \tilde{J}^-, \quad \tilde{H} = -\frac{2}{\ell} \tilde{J}^-, \\ \tilde{S} = \tilde{S}^+ + \tilde{S}^-, \quad \tilde{M} = -\frac{2}{\ell} \tilde{S}^-. \tag{3.9}$$

Although the central charges can be set to zero, their presence is essential in the torsional NR algebra (3.7) in order to admit a non-degenerate invariant bilinear trace:

$$\langle \tilde{J} \tilde{S} \rangle = -\tilde{\alpha}_0, \\ \langle \tilde{G}_a \tilde{G}_b \rangle = \tilde{\alpha}_0 \delta_{ab}, \\ \langle \tilde{G}_a \tilde{P}_b \rangle = \tilde{\alpha}_1 \delta_{ab}, \\ \langle \tilde{H} \tilde{S} \rangle = \langle \tilde{M} \tilde{J} \rangle = -\tilde{\alpha}_1, \\ \langle \tilde{P}_a \tilde{P}_b \rangle = -\frac{2\tilde{\alpha}_1}{\ell} \delta_{ab}, \\ \langle \tilde{H} \tilde{M} \rangle = \frac{2\tilde{\alpha}_1}{\ell}, \tag{3.10}$$

where the relativistic parameters α 's have been rescaled as

$$\alpha_0 = \tilde{\alpha}_0 \xi^2, \quad \alpha_1 = \tilde{\alpha}_1 \xi. \tag{3.11}$$

Analogously to [74,79,102], such rescaling is required in order to have a finite NR CS action after the IW contraction. One can see that the non-degeneracy of the invariant tensor is preserved as long as $\tilde{\alpha}_0 \neq -\frac{\ell}{2} \tilde{\alpha}_1$ and $\tilde{\alpha}_1 \neq 0$. The non-degeneracy of the invariant tensor is related to the requirement that the NR CS gravity action involves a kinematical term for each gauge field. In particular, the corresponding gauge connection one-form for the torsional NR algebra (3.7) reads

$$\tilde{A} = \omega \tilde{J} + \omega^a \tilde{G}_a + \tau \tilde{H} + e^a \tilde{P}_a + m \tilde{M} + s \tilde{S}. \tag{3.12}$$

The curvature associated to this gauge connection is given by

$$\tilde{F} = R(\omega)\tilde{J} + R^a(\omega^b)\tilde{G}_a + R(\tau)\tilde{H} + R^a(e^b)\tilde{P}_a + R(m)\tilde{M} + R(s)\tilde{S}, \tag{3.13}$$

where

$$\begin{aligned} R(\omega) &= d\omega, \\ R^a(\omega^b) &= d\omega^a + \epsilon^{ac}\omega\omega_c, \\ R(\tau) &= d\tau, \\ R^a(e^b) &= de^a + \epsilon^{ac}\omega e_c + \epsilon^{ac}\tau\omega_c - \frac{2}{\ell}\epsilon^{ac}\tau e_c, \\ R(m) &= dm + \epsilon^{ac}e_a\omega_c - \frac{1}{\ell}\epsilon^{ac}e_a e_c, \\ R(s) &= ds + \frac{1}{2}\epsilon^{ac}\omega_a\omega_c. \end{aligned} \tag{3.14}$$

The respective CS action based on the torsional NR algebra (3.7) is obtained considering the non-vanishing components of the invariant tensor (3.10) and the gauge connection one-form (3.12) in the general expression of the CS action (2.6):

$$\begin{aligned} I_{\text{NR}} &= \frac{k}{4\pi} \int \tilde{\alpha}_0 \left[\omega_a R^a(\omega^b) - 2sR(\omega) \right] \\ &\quad + \tilde{\alpha}_1 \left[2e_a R^a(\omega^b) - 2mR(\omega) - 2\tau R(s) \right. \\ &\quad \left. - \frac{2}{\ell}e_a R^a(e^b) + \frac{2}{\ell}mR(\tau) + \frac{2}{\ell}\tau R(m) + \frac{2}{\ell^2}\tau\epsilon^{ac}e_a e_c \right] \end{aligned} \tag{3.15}$$

The NR CS action (3.15) is the most general NR CS action invariant under the torsional NR algebra (3.7). One can notice that it contains two independent sectors and can be seen as the NR version of the three-dimensional teleparallel CS gravity theory presented in [17]. The first term is indistinguishable from the exotic sector of the extended Bargmann gravity [57], being the corresponding NR version of the exotic gravity [89]. The term proportional to $\tilde{\alpha}_1$ can be seen as a cosmological extension of the extended Bargmann gravity action [54]. However, the inclusion of the cosmological constant through the torsional NR algebra is quite different from the one constructed using the extended Newton–Hooke symmetry. Indeed, the field equations derived from (3.15) are given by the vanishing of the curvature two-forms (3.14). One can see that the field equations for ω^a , e^a , s and m are diverse from the Newton–Hooke ones, the latter being given by

$$\begin{aligned} F^a(\omega^b) &= d\omega^a + \epsilon^{ac}\omega\omega_c + \frac{1}{\ell^2}\epsilon^{ac}\tau e_c = 0, \\ F^a(e^b) &= de^a + \epsilon^{ac}\omega e_c + \epsilon^{ac}\tau\omega_c = 0, \\ F(s) &= ds + \frac{1}{2}\epsilon^{ac}\omega_a\omega_c + \frac{1}{2\ell^2}\epsilon^{ac}e_a e_c = 0, \\ F(m) &= dm + \epsilon^{ac}e_a\omega_c = 0. \end{aligned} \tag{3.16}$$

In our torsional NR gravity theory, the cosmological constant can be seen as a source for the spatial torsion $T^a(e^b) = de^a + \epsilon^{ac}\omega e_c + \epsilon^{ac}\tau\omega_c$ and for the curvature $T(m) = dm + \epsilon^{ac}e_a\omega_c$. Indeed, on-shell we find

$$\begin{aligned} T^a(e^b) &= \frac{2}{\ell}\epsilon^{ac}\tau e_c, \\ T(m) &= \frac{1}{\ell}\epsilon^{ac}e_a e_c. \end{aligned} \tag{3.17}$$

The NR counterpart of the teleparallel CS gravity [17] is then described, on-shell, by vanishing curvatures $R(\omega) = 0 = R^a(\omega^a)$ and a non-vanishing spatial torsion $T^a(e^b) \neq 0$. Naturally, we recover the extended Bargmann gravity theory [54] with a vanishing spatial torsion in the vanishing cosmological constant $\ell \rightarrow \infty$.

Let us note that the NR gravity action (3.15) can be alternatively recovered from the relativistic $U(1)$ -enlarged teleparallel CS action (3.3). Indeed, one can express the relativistic gauge fields in terms of the NR ones as follows:

$$\begin{aligned} W^0 &= \omega + \frac{s}{2\xi^2}, & W^a &= \frac{\omega^a}{\xi}, & S &= \omega - \frac{s}{2\xi^2}, \\ E^0 &= \xi\tau + \frac{m}{2\xi}, & E^a &= e^a, & M &= \xi\tau - \frac{m}{2\xi}, \end{aligned} \tag{3.18}$$

which satisfy $A = \tilde{A}$. The NR CS action (3.15) appears considering these last expressions along the rescaling of the relativistic parameters (3.11) on the relativistic CS action (3.3) and then applying the limit $\xi \rightarrow \infty$.

The presence of torsion in a NR environment has already been explored in [53, 83, 84] describing a (twistless) torsional Newton–Cartan gravity. However, unlike the present case, the torsional Newton–Cartan gravity is described by a non-vanishing time-like torsion. Such torsional model can be obtained by gauging the Schrödinger algebra which corresponds to the conformal extension of the Bargmann algebra [85]. A non-zero torsion condition in a NR framework was first encountered in the context of Lifshitz holography [25] and Quantum Hall Effect [34]. Although the NR version of the teleparallel gravity contains a zero time-like torsion, the presence of a non-vanishing spatial torsion can be viewed as a step towards introducing non-zero time-like torsion as in the (twistless) torsional Newton–Cartan gravity. In particular, one can note that under boosts (with parameter λ^a), one has that

$$\begin{aligned} \delta T^a(e) &= \frac{2}{\ell}\epsilon^{ac}\lambda_a R(\tau), \\ \delta T(m) &= \frac{1}{\ell}\epsilon^{ac}\lambda_a T_c(e), \end{aligned} \tag{3.19}$$

which is consistent with the field equation for the torsion (3.17). This boost behaviour also implies that introducing non-zero time-like torsion $R(\tau)$ requires introducing non-zero $T^a(e)$ and $T(m)$. Since the CS formulation offers us a simpler way to construct the respective supersymmetric

extension, our construction could be then useful to approach a NR supergravity model with a non-zero time-like supertorsion.

4 Non-relativistic supergravity with torsion

We will now present the explicit supersymmetric extension of the torsional NR gravity theory previously discussed. The NR action will be characterized, on-shell, by the vanishing of the NR curvatures but non-vanishing spatial super-torsion.

A torsional NR superalgebra can be obtained by applying the S-expansion method to the $\mathcal{N} = 2$ teleparallel superalgebra. The S-expansion procedure will also provide us with a non-degenerate invariant tensor and therefore we will be able to produce a well-defined torsional NR CS supergravity action which can be seen as the corresponding NR counterpart of the $\mathcal{N} = 2$ teleparallel supergravity action presented in [17].

4.1 Torsional non-relativistic superalgebra

We shall now move on to the supersymmetric extension of the results presented in the previous section. To this aim, we will focus on the relativistic $\mathcal{N} = 2$ teleparallel superalgebra presented in [17], and study its non-relativistic expansion.

At the relativistic level, \mathcal{N} -extended teleparallel supergravity has been formulated as a CS theory invariant under the so-called \mathcal{N} -extended teleparallel superalgebra [17]. Although both teleparallel and AdS CS supergravity theories correspond to cosmological extensions of the super-Poincaré one, they are described by different equations of motion. Indeed, the key difference between both CS models resides in the fact that, unlike the AdS supergravity, in the teleparallel case we have, on-shell, a non-vanishing supertorsion, while the Lorentz supercurvature vanishes (cf. also [103] for a detailed analysis of three-dimensional supergravity).

In the $\mathcal{N} = 2$ case, which is the one of interest in the present paper, the teleparallel superalgebra is spanned by the bosonic generators (J_A, P_A) along with the internal symmetry generator Z , and the fermionic charges Q_α^i , with $i = 1, 2$ and $\alpha = 1, 2$, which are two-components Majorana spinor charges.³ However, in order to have a well-defined flat limit $\ell \rightarrow \infty$ leading to the $\mathcal{N} = 2$ Poincaré superalgebra extended with an $\mathfrak{so}(2)$ automorphism algebra [104], the $\mathcal{N} = 2$ teleparallel superalgebra was extended in [17] to include the automorphism generator S . To this end, the redefinition $\mathcal{T} = Z - \frac{\ell}{2}S$ was performed, obtaining the following

non-trivial (anti-)commutation relations in the $\mathcal{N} = 2$ case:

$$\begin{aligned} [J_A, Q_\alpha^i] &= -\frac{1}{2} (\gamma_A)_\alpha^\beta Q_\beta^i, \\ [\mathcal{T}, Q_\alpha^i] &= -\epsilon^{ij} Q_\alpha^j, \\ \{Q_\alpha^i, Q_\beta^j\} &= -\delta^{ij} (\gamma^A C)_{\alpha\beta} \left(P_A + \frac{2}{\ell} J_A \right) \\ &\quad + C_{\alpha\beta} \epsilon^{ij} \left(\frac{2}{\ell} \mathcal{T} + S \right), \end{aligned} \tag{4.1}$$

along with (2.1). The gamma matrices in three dimensions are denoted by γ_A and C is the charge conjugation matrix, satisfying $C^T = -C$ and $C\gamma^A = (C\gamma^A)^T$, while ϵ^{ij} is the rank-2 Levi-Civita symbol ($\epsilon^{12} = 1, \epsilon^{21} = -1$). The $\mathcal{N} = 2$ Poincaré superalgebra extended with the $\mathfrak{so}(2)$ automorphism algebra is properly recovered in the flat limit $\ell \rightarrow \infty$. The presence of the automorphism generator in the Poincaré case is required in order to define a non-degenerate invariant tensor [104]. Observe that, as in the case of the $\mathcal{N} = 2$ Poincaré superalgebra, the generator S is a central charge. Even though the $\mathcal{N} = 2$ teleparallel superalgebra given by (2.1) and (4.1) presents a well-defined Poincaré limit, its (anti-)commutation relations are quite different from the super-AdS ones. Let us stress that the $\mathcal{N} = 2$ teleparallel superalgebra can be written as the direct sum of $\mathfrak{osp}(2|2) \otimes \mathfrak{sp}(2)$ and the $\mathfrak{so}(2)$ automorphism algebra by considering the following identification of the generators:

$$\begin{aligned} L_A &\equiv J_A + \frac{\ell}{2} P_A, \quad S_A \equiv -\frac{\ell}{2} P_A, \\ \mathcal{G}_\alpha^i &\equiv \sqrt{\frac{\ell}{2}} Q_\alpha^i, \quad \mathcal{M} \equiv \mathcal{T} + \frac{\ell}{2} S, \quad B \equiv -\frac{\ell}{2} S, \end{aligned} \tag{4.2}$$

where $(L_A, S_A, \mathcal{M}, \mathcal{G}_\alpha^i)$ satisfy the $\mathfrak{osp}(2|2) \otimes \mathfrak{sp}(2)$ superalgebra, while B is the $\mathfrak{so}(2)$ automorphism generator.

The $\mathcal{N} = 2$ teleparallel superalgebra is endowed with the following non-vanishing components of the (non-degenerate) invariant tensor:

$$\begin{aligned} \langle \mathcal{T}\mathcal{T} \rangle &= -2\alpha_0, \\ \langle \mathcal{T}\mathcal{S} \rangle &= -2\alpha_1, \\ \langle \mathcal{S}\mathcal{S} \rangle &= \frac{4\alpha_1}{\ell}, \\ \langle Q_\alpha^i Q_\beta^j \rangle &= 2 \left(\frac{2\alpha_0}{\ell} + \alpha_1 \right) C_{\alpha\beta} \delta^{ij}, \end{aligned} \tag{4.3}$$

along with (2.2). We refer the reader to [17] for the way in which the constant parameters α_0 and α_1 are related to the $\mathfrak{osp}(2|2) \otimes \mathfrak{sp}(2)$ and $\mathfrak{so}(2)$ ones. The flat limit $\ell \rightarrow \infty$ yields the invariant tensor for the $\mathcal{N} = 2$ Poincaré superalgebra extended with $\mathfrak{so}(2)$ [104].

A NR teleparallel superalgebra can be obtained from the $\mathcal{N} = 2$ relativistic one by performing an S-expansion [60]. Indeed, as it was also seen in [64–66], the S-expansion of a relativistic superalgebra considering $S_E^{(2)}$ as the relevant

³ For simplicity, in the following we will frequently omit writing of the spinor indices.

semigroup reproduces a NR contraction. This is due to the fact that the S-expansion method can be seen as a generalization of the Inönö–Wigner contraction process [99] when the semigroup considered belongs to the $S_E^{(N)}$ family. Here we focus on the semigroup $S_E^{(2)}$, which allows to obtain not only a well-defined NR teleparallel superalgebra but also its non-degenerate invariant tensor.

We first proceed as in Sect. 3 by considering the decomposition of the A index given in (3.4). Besides, we perform a redefinition of the supercharges by defining

$$Q_\alpha^\pm = \frac{1}{\sqrt{2}} \left(Q_\alpha^1 \pm (\gamma^0)_{\alpha\beta} Q_\beta^2 \right). \tag{4.4}$$

Then, we consider $S_E^{(2)} = \{\lambda_0, \lambda_1, \lambda_2, \lambda_3\}$ as the relevant semigroup to perform the expansion of the relativistic $\mathcal{N} = 2$ teleparallel superalgebra. The elements of $S_E^{(2)}$ satisfy the following multiplication law:

λ_3	$\lambda_3 \lambda_3 \lambda_3 \lambda_3$	(4.5)
λ_2	$\lambda_2 \lambda_3 \lambda_3 \lambda_3$	
λ_1	$\lambda_1 \lambda_2 \lambda_3 \lambda_3$	
λ_0	$\lambda_0 \lambda_1 \lambda_2 \lambda_3$	
	$\lambda_0 \lambda_1 \lambda_2 \lambda_3$	

where $\lambda_3 = 0_s$ is the zero element of the semigroup such that $0_s \lambda_k = 0_s, k = 1, 2, 3$. Before applying the S-expansion procedure, we consider a particular subspace decomposition of the $\mathcal{N} = 2$ teleparallel superalgebra. To this end, let us first consider the split of the index A along with the redefinition of the supercharges above. Hence, we can see that the subspaces decomposition of the teleparallel superalgebra given by $V_0 = \{J_0, P_0, T, S, Q_\alpha^+\}$ and $V_1 = \{J_a, P_a, Q_\alpha^-\}$ satisfies

$$[V_0, V_0] \subset V_0, \quad [V_0, V_1] \subset V_1, \quad [V_1, V_1] \subset V_0. \tag{4.6}$$

Let us consider now $S_E^{(2)} = S_0 \cup S_1$ as decomposition of the relevant semigroup $S_E^{(2)}$, where

$$\begin{aligned} S_0 &= \{\lambda_0, \lambda_2, \lambda_3\}, \\ S_1 &= \{\lambda_1, \lambda_3\}. \end{aligned} \tag{4.7}$$

Then, the decomposition (4.7) is said to be resonant since it satisfies the same structure as the subspaces, that is

$$S_0 \cdot S_0 \subset S_0, \quad S_0 \cdot S_1 \subset S_1, \quad S_1 \cdot S_1 \subset S_0. \tag{4.8}$$

Following [60], after extracting a resonant subalgebra of the $S_E^{(2)}$ -expansion of the $\mathcal{N} = 2$ teleparallel superalgebra and applying a 0_s -reduction, one ends up with a new NR expanded superalgebra spanned by the set of generators

$$\{\tilde{J}, \tilde{G}_a, \tilde{S}, \tilde{H}, \tilde{P}_a, \tilde{M}, \tilde{T}_1, \tilde{T}_2, \tilde{U}_1, \tilde{U}_2, \tilde{Q}_\alpha^+, \tilde{R}_\alpha, \tilde{Q}_\alpha^-\}, \tag{4.9}$$

which are related to the relativistic ones through the semigroup elements as

λ_3		
λ_2	$\tilde{S}, \tilde{M}, \tilde{T}_2, \tilde{U}_2, \tilde{R}_\alpha$	
λ_1		$\tilde{G}_a, \tilde{P}_a, \tilde{Q}_\alpha^-$
λ_0	$\tilde{J}, \tilde{H}, \tilde{T}_1, \tilde{U}_1, \tilde{Q}_\alpha^+$	
	$J_0, P_0, T, S, Q_\alpha^+$	J_a, P_a, Q_α^-

(4.10)

The NR generators satisfy precisely the purely bosonic subalgebra (3.7) along with the following (anti-)commutation relations:

$$\begin{aligned} [\tilde{J}, \tilde{Q}_\alpha^\pm] &= -\frac{1}{2} (\gamma_0)_\alpha^\beta \tilde{Q}_\beta^\pm, & [\tilde{J}, \tilde{R}_\alpha] &= -\frac{1}{2} (\gamma_0)_\alpha^\beta \tilde{R}_\beta, \\ [\tilde{S}, \tilde{Q}_\alpha^+] &= -\frac{1}{2} (\gamma_0)_\alpha^\beta \tilde{R}_\beta, \\ [\tilde{G}_a, \tilde{Q}_\alpha^+] &= -\frac{1}{2} (\gamma_a)_\alpha^\beta \tilde{Q}_\beta^-, & [\tilde{G}_a, \tilde{Q}_\alpha^-] &= -\frac{1}{2} (\gamma_a)_\alpha^\beta \tilde{R}_\beta, \\ [\tilde{T}_1, \tilde{Q}_\alpha^\pm] &= \pm (\gamma^0)_{\alpha\beta} \tilde{Q}_\beta^\pm, & [\tilde{T}_2, \tilde{Q}_\alpha^+] &= (\gamma^0)_{\alpha\beta} \tilde{R}_\beta, \\ [\tilde{T}_1, \tilde{R}_\alpha] &= (\gamma^0)_{\alpha\beta} \tilde{R}_\beta, \\ \{\tilde{Q}_\alpha^+, \tilde{Q}_\beta^+\} &= -(\gamma^0 C)_{\alpha\beta} \left(\tilde{H} + \frac{2}{\ell} \tilde{J} \right) - (\gamma^0 C)_{\alpha\beta} \left(\frac{2}{\ell} \tilde{T}_1 + \tilde{U}_1 \right), \\ \{\tilde{Q}_\alpha^-, \tilde{Q}_\beta^-\} &= -(\gamma^0 C)_{\alpha\beta} \left(\tilde{M} + \frac{2}{\ell} \tilde{S} \right) + (\gamma^0 C)_{\alpha\beta} \left(\frac{2}{\ell} \tilde{T}_2 + \tilde{U}_2 \right), \\ \{\tilde{Q}_\alpha^+, \tilde{R}_\beta\} &= -(\gamma^0 C)_{\alpha\beta} \left(\tilde{M} + \frac{2}{\ell} \tilde{S} \right) - (\gamma^0 C)_{\alpha\beta} \left(\frac{2}{\ell} \tilde{T}_2 + \tilde{U}_2 \right), \\ \{\tilde{Q}_\alpha^+, \tilde{Q}_\beta^-\} &= -(\gamma^a C)_{\alpha\beta} \left(\tilde{P}_a + \frac{2}{\ell} \tilde{G}_a \right). \end{aligned} \tag{4.11}$$

The superalgebra given by (3.7) and (4.11) can be seen as the supersymmetric extension of the torsional NR algebra obtained previously (3.7). One can see that such superalgebra requires the introduction of four additional bosonic generators $\tilde{T}_1, \tilde{T}_2, \tilde{U}_1$ and \tilde{U}_2 . Notice that the generators \tilde{U}_1 and \tilde{U}_2 are central charges, while \tilde{T}_1 and \tilde{T}_2 which arise from the S-expansion of the relativistic R-symmetry generator T , act non-trivially on the spinor charges \tilde{Q}_α^\pm and \tilde{R}_α . Although the central charges can be set to zero by performing a simple contraction, their presence here is essential in order to have a non-degenerate invariant tensor. Interestingly, we recover the extended Bargmann superalgebra [54] in the flat limit $\ell \rightarrow \infty$ and considering $\tilde{T}_1 = \tilde{T}_2 = \tilde{U}_1 = \tilde{U}_2 = 0$. Such behavior is inherited from the respective relativistic counterparts, in which case the teleparallel algebra reproduces the Poincaré algebra in the vanishing cosmological constant limit. The inclusion of a cosmological constant into the extended Bargmann supergravity has already been approached considering a supersymmetric extension of the extended Newton–Hooke symmetry [57]. Here we present an alternative NR superalgebra allowing the presence of a cosmological constant, but with a non-vanishing spatial super-torsion as we shall see.

Before approaching the construction of a CS supergravity action based on the torsional NR superalgebra, let us stress

that the NR superalgebra given by (3.7) and (4.11) can be written as two copies of the Nappi–Witten algebra, one of which is augmented by supersymmetry endowed with $u(1)$ generators. To this end, let us consider the following redefinition of the generators of the torsional NR superalgebra:

$$\begin{aligned} \tilde{G}_a &= G_a + \hat{G}_a, \quad \tilde{P}_a = -\frac{2}{\ell}\hat{G}_a, \quad \tilde{Q}_\alpha^+ = \sqrt{\frac{2}{\ell}}Q_\alpha^+ \\ \tilde{J} &= J + \hat{J}, \quad \tilde{H} = -\frac{2}{\ell}\hat{J}, \quad \tilde{Q}_\alpha^- = \sqrt{\frac{2}{\ell}}Q_\alpha^-, \\ \tilde{S} &= S + \hat{S}, \quad \tilde{M} = -\frac{2}{\ell}\hat{S}, \quad \tilde{R}_\alpha = \sqrt{\frac{2}{\ell}}\mathcal{R}_\alpha, \\ \tilde{T}_1 &= X_1 + \hat{X}_1, \quad \tilde{U}_1 = -\frac{2}{\ell}\hat{X}_1, \\ \tilde{T}_2 &= X_2 + \hat{X}_2, \quad \tilde{U}_2 = -\frac{2}{\ell}\hat{X}_2. \end{aligned} \tag{4.12}$$

Here, the subset spanned by $\{J, G_a, S, X_1, X_2, Q_\alpha^+, Q_\alpha^-, \mathcal{R}_\alpha\}$ satisfy a supersymmetric extension of the Nappi–Witten algebra [57,65]:

$$\begin{aligned} [J, G_a] &= \epsilon_{ab}G_b, & [G_a, G_b] &= -\epsilon_{ab}S, \\ [J, Q_\alpha^\pm] &= -\frac{1}{2}(\gamma_0)_\alpha^\beta Q_\beta^\pm, & [J, \mathcal{R}_\alpha] &= -\frac{1}{2}(\gamma_0)_\alpha^\beta \mathcal{R}_\beta, \\ [G_a, Q_\alpha^+] &= -\frac{1}{2}(\gamma_a)_\alpha^\beta Q_\beta^-, & [G_a, Q_\alpha^-] &= -\frac{1}{2}(\gamma_a)_\alpha^\beta \mathcal{R}_\beta, \\ [S, Q_\alpha^+] &= -\frac{1}{2}(\gamma_0)_\alpha^\beta \mathcal{R}_\beta, & [X_1, Q_\alpha^\pm] &= \pm\frac{1}{2}(\gamma_0)_{\alpha\beta} Q_\beta^\pm, \\ [X_2, Q_\alpha^+] &= \frac{1}{2}(\gamma_0)_{\alpha\beta} \mathcal{R}_\beta, & [X_1, \mathcal{R}_\alpha] &= \frac{1}{2}(\gamma_0)_{\alpha\beta} \mathcal{R}_\beta, \\ \{Q_\alpha^+, Q_\beta^-\} &= -(\gamma^a C)_{\alpha\beta} G_a, \\ \{Q_\alpha^+, Q_\beta^+\} &= -(\gamma^0 C)_{\alpha\beta} J - (\gamma^0 C)_{\alpha\beta} X_1, \\ \{Q_\alpha^-, Q_\beta^-\} &= -(\gamma^0 C)_{\alpha\beta} S + (\gamma^0 C)_{\alpha\beta} X_2, \\ \{Q_\alpha^+, \mathcal{R}_\beta\} &= -(\gamma^0 C)_{\alpha\beta} S - (\gamma^0 C)_{\alpha\beta} X_2. \end{aligned} \tag{4.13}$$

On the other hand, the set of generators $\{\tilde{J}, \tilde{G}_a, \tilde{S}, \tilde{X}_1, \tilde{X}_2\}$ satisfy a $U(1)$ -enlargement of the usual Nappi–Witten algebra (3.8).

4.2 Torsional non-relativistic supergravity theory

Let us construct the explicit NR CS supergravity action invariant under the torsional NR superalgebra (3.7) and (4.11). Although the Nappi–Witten algebraic structure seems simpler, we are interested in exploring an alternative way to accommodate a cosmological constant into the extended Bargmann supergravity different from considering the extended Newton–Hooke one [63]. As we shall see, the inclusion of a cosmological constant through the NR superalgebra (3.7) and (4.11) will imply the presence of a non-vanishing spatial super-torsion.

One can show that the non-vanishing components of an invariant tensor for the NR teleparallel superalgebra just obtained are given by (3.10) along with

$$\begin{aligned} \langle \tilde{T}_1 \tilde{T}_2 \rangle &= -2\tilde{\alpha}_0, \\ \langle \tilde{T}_1 \tilde{U}_2 \rangle &= \langle \tilde{T}_2 \tilde{U}_1 \rangle = -2\tilde{\alpha}_1, \\ \langle \tilde{U}_1 \tilde{U}_2 \rangle &= \frac{4\tilde{\alpha}_1}{\ell}, \\ \langle Q_\alpha^+ R_\beta \rangle &= 2\left(\frac{2\tilde{\alpha}_0}{\ell} + \tilde{\alpha}_1\right) C_{\alpha\beta}, \\ \langle Q_\alpha^- Q_\beta^- \rangle &= 2\left(\frac{2\tilde{\alpha}_0}{\ell} + \tilde{\alpha}_1\right) C_{\alpha\beta}. \end{aligned} \tag{4.14}$$

Here the invariant tensor can be obtained from the relativistic one (2.2) and (4.3) considering (4.10) and defining the NR parameters in terms of the relativistic ones through the semigroup elements as

$$\tilde{\alpha}_0 = \lambda_2 \alpha_0, \quad \tilde{\alpha}_1 = \lambda_2 \alpha_1. \tag{4.15}$$

The gauge connection one-form for the torsional NR superalgebra given by (3.7) and (4.14) reads as follows:

$$\begin{aligned} \tilde{A} &= \omega \tilde{J} + \omega^a \tilde{G}_a + \tau \tilde{H} + e^a \tilde{P}_a + m \tilde{M} + s \tilde{S} + t_1 \tilde{T}_1 \\ &\quad + t_2 \tilde{T}_2 + u_1 \tilde{U}_1 + u_2 \tilde{U}_2 + \tilde{\psi}^+ \tilde{Q}^+ + \tilde{\psi}^- \tilde{Q}^- + \tilde{\rho} \tilde{R}. \end{aligned} \tag{4.16}$$

The curvature two-form associated with the gauge connection above is given by

$$\begin{aligned} \tilde{F} &= F(\omega) \tilde{J} + F^a(\omega^b) \tilde{G}_a + F(\tau) \tilde{H} \\ &\quad + F^a(e^b) \tilde{P}_a + F(m) \tilde{M} + F(s) \tilde{S} \\ &\quad + F(t_1) \tilde{T}_1 + F(t_2) \tilde{T}_2 + F(u_1) \tilde{U}_1 + F(u_2) \tilde{U}_2 \\ &\quad + \nabla \tilde{\psi}^+ \tilde{Q}^+ + \nabla \tilde{\psi}^- \tilde{Q}^- + \nabla \tilde{\rho} \tilde{R}, \end{aligned} \tag{4.17}$$

where the bosonic curvatures are given by

$$\begin{aligned} F(\omega) &= R(\omega) + \frac{1}{\ell} \tilde{\psi}^+ \gamma^0 \psi^+ = d\omega + \frac{1}{\ell} \tilde{\psi}^+ \gamma^0 \psi^+, \\ F^a(\omega^b) &= R^a(\omega^b) + \frac{2}{\ell} \tilde{\psi}^+ \gamma^a \psi^- = d\omega^a + \epsilon^{ac} \omega \omega_c + \frac{2}{\ell} \tilde{\psi}^+ \gamma^a \psi^-, \\ F(\tau) &= R(\tau) + \frac{1}{2} \tilde{\psi}^+ \gamma^0 \psi^+ = d\tau + \frac{1}{2} \tilde{\psi}^+ \gamma^0 \psi^+, \\ F^a(e^b) &= R^a(e^b) + \tilde{\psi}^+ \gamma^a \psi^- = de^a + \epsilon^{ac} \omega e_c + \epsilon^{ac} \tau \omega_c \\ &\quad - \frac{2}{\ell} \epsilon^{ac} \tau e_c + \tilde{\psi}^+ \gamma^a \psi^-, \\ F(m) &= R(m) + \frac{1}{2} \tilde{\psi}^- \gamma^0 \psi^- + \tilde{\psi}^+ \gamma^0 \rho \\ &= dm + \epsilon^{ac} e_a \omega_c - \frac{1}{\ell} \epsilon^{ac} e_a e_c + \frac{1}{2} \tilde{\psi}^- \gamma^0 \psi^- + \tilde{\psi}^+ \gamma^0 \rho, \\ F(s) &= R(s) + \frac{1}{\ell} \tilde{\psi}^- \gamma^0 \psi^- + \frac{2}{\ell} \tilde{\psi}^+ \gamma^0 \rho \\ &= ds + \frac{1}{2} \epsilon^{ac} \omega_a \omega_c + \frac{1}{\ell} \tilde{\psi}^- \gamma^0 \psi^- + \frac{2}{\ell} \tilde{\psi}^+ \gamma^0 \rho, \\ F(t_1) &= dt_1 + \frac{1}{\ell} \tilde{\psi}^+ \gamma^0 \psi^+, \\ F(t_2) &= dt_2 - \frac{1}{\ell} \tilde{\psi}^- \gamma^0 \psi^- + \frac{2}{\ell} \tilde{\psi}^+ \gamma^0 \rho, \end{aligned}$$

$$\begin{aligned}
 F(u_1) &= du_1 + \frac{1}{2}\bar{\psi}^+\gamma^0\psi^+, \\
 F(u_2) &= du_2 - \frac{1}{2}\bar{\psi}^-\gamma^0\psi^- + \bar{\psi}^+\gamma^0\rho,
 \end{aligned}
 \tag{4.18}$$

with $R(\omega)$, $R^a(\omega^b)$, $R(\tau)$, $R^a(e^b)$, $R(m)$, and $R(s)$ defined in (3.14), while the fermionic field-strengths are

$$\begin{aligned}
 \nabla\psi^+ &= d\psi^+ + \frac{1}{2}\omega\gamma_0\psi^+ + t_1\gamma_0\psi^+, \\
 \nabla\psi^- &= d\psi^- + \frac{1}{2}\omega\gamma_0\psi^- + \frac{1}{2}\omega^a\gamma_a\psi^+ - t_1\gamma_0\psi^-, \\
 \nabla\rho &= d\rho + \frac{1}{2}\omega\gamma_0\rho + \frac{1}{2}s\gamma_0\psi^+ + \frac{1}{2}\omega^a\gamma_a\psi^- \\
 &\quad + t_2\gamma_0\psi^+ + t_1\gamma_0\rho.
 \end{aligned}
 \tag{4.19}$$

The supergravity CS action based on the torsional NR superalgebra given by (3.7) and (4.14) is then obtained by considering the non-vanishing components of the invariant tensor (3.10) and (4.14) together with the gauge connection one-form (4.16) in the general expression of the CS action (2.6). By doing so, we get

$$\begin{aligned}
 I_{\text{NR}}^{\text{super}} &= \frac{k}{4\pi} \int \tilde{\alpha}_0 \left[\omega_a R^a(\omega^b) - 2sR(\omega) - 4t_1 dt_2 \right. \\
 &\quad \left. - \frac{4}{\ell} \bar{\psi}^+ \nabla \rho - \frac{4}{\ell} \bar{\rho} \nabla \psi^+ - \frac{4}{\ell} \bar{\psi}^- \nabla \psi^- \right] \\
 &\quad + 2\tilde{\alpha}_1 \left[e_a R^a(\omega^b) - mR(\omega) - \tau R(s) - \frac{1}{\ell} e_a R^a(e^b) \right. \\
 &\quad \left. + \frac{1}{\ell} mR(\tau) + \frac{1}{\ell} \tau R(m) \right. \\
 &\quad \left. + \frac{1}{\ell^2} \tau \epsilon^{ac} e_a e_c - 2t_1 du_2 - 2t_2 du_1 \right. \\
 &\quad \left. + \frac{4}{\ell} u_1 du_2 - \bar{\psi}^+ \nabla \rho - \bar{\rho} \nabla \psi^+ - \bar{\psi}^- \nabla \psi^- \right].
 \end{aligned}
 \tag{4.20}$$

The NR CS supergravity action (4.20) contains two independent sectors and can be seen as the NR version of the $\mathcal{N} = 2$ relativistic teleparallel CS supergravity action. The first term corresponds to a supersymmetric NR exotic action, while the contribution proportional to $\tilde{\alpha}_1$ can be seen as a supersymmetric cosmological extension of the extended Bargmann gravity action [54]. In the vanishing cosmological constant limit $\ell \rightarrow \infty$ we recover the most general extended Bargmann supergravity action [57]. In particular, the term proportional to $\tilde{\alpha}_0$ is no more supersymmetric in the flat limit. Neglecting the fermionic contributions and the additional bosonic gauge fields $\{t_1, t_2, u_1, u_2\}$, (4.20) precisely boils down to the purely NR bosonic action (3.15) we have previously presented.

Let us stress that the non-degeneracy of the invariant tensor is preserved as long as $\tilde{\alpha}_0 \neq -\frac{\ell}{2}\tilde{\alpha}_1$ and $\tilde{\alpha}_1 \neq 0$, which implies that the field equations derived from (4.20) give the vanishing of the curvatures (4.18) and (4.19). Here, analogously to the purely bosonic case, the cosmological constant can be seen as a source for the spatial super-torsion $\hat{T}^a(e^b) = de^a + \epsilon^{ac}\omega e_c + \epsilon^{ac}\tau e_c + \bar{\psi}^+\gamma^a\psi^-$ and for the

curvature $\hat{T}(m) = dm + \epsilon^{ac}e_a\omega_c + \frac{1}{2}\bar{\psi}^-\gamma^0\psi^- + \bar{\psi}^+\gamma^0\rho$. In particular, on-shell we have

$$\begin{aligned}
 \hat{T}^a(e^b) &= \frac{2}{\ell}\epsilon^{ac}\tau e_c, \\
 \hat{T}(m) &= \frac{1}{\ell}\epsilon^{ac}e_a e_c.
 \end{aligned}
 \tag{4.21}$$

Hence, the torsional NR supergravity theory is then characterized, on-shell, by vanishing supercurvatures $F(\omega) = 0 = F^a(\omega^b)$ and a non-vanishing spatial super-torsion $\hat{T}^a(e^b) \neq 0$. In the vanishing cosmological constant limit $\ell \rightarrow \infty$, the field equations reduces to those of the extended Bargmann ones in which the super-torsion vanishes. Let us note that the present theory is quite different from the extended Newton–Hooke supergravity at the dynamic level. Indeed, the Newton–Hooke supergravity theory is characterized by a vanishing super-torsion and the cosmological constant appears explicitly along the supercurvatures of the ω and ω^a NR gauge fields [57,63].

As an ending remark, one can alternatively obtain the torsional NR supergravity action (4.20) from the relativistic $\mathcal{N} = 2$ teleparallel CS supergravity action [17]:

$$\begin{aligned}
 I_{\text{TSG}}^{\mathcal{N}=2} &= \frac{k}{4\pi} \int_{\mathcal{M}} \left\{ \alpha_0 \left(W^A dW_A \right. \right. \\
 &\quad \left. \left. + \frac{1}{3}\epsilon^{ABC} W_A W_B W_C - 2Ad\mathcal{A} - \frac{4}{\ell} \bar{\Psi}^i \nabla \Psi^i \right) \right. \\
 &\quad \left. + \alpha_1 \left(2E_A R^A(W) + \frac{4}{3\ell^2} \epsilon^{ABC} E_A E_B E_C \right. \right. \\
 &\quad \left. \left. - \frac{2}{\ell} T^A E_A - 4Cd\mathcal{A} + \frac{4}{\ell} Cd\mathcal{C} - 2\bar{\Psi}^i \nabla \Psi^i \right) \right\},
 \end{aligned}
 \tag{4.22}$$

where $i = 1, 2$ and

$$\begin{aligned}
 R^A(W) &= dW^A + \frac{1}{2}\epsilon^{ABC} W_B W_C, \\
 T^A &= dE^A + \epsilon^{ABC} W_B E_C, \\
 \nabla \Psi^i &= d\Psi^i + \frac{1}{2}W^A \gamma_A \Psi^i + \epsilon^{ij} \mathcal{A} \Psi^j.
 \end{aligned}
 \tag{4.23}$$

The NR supergravity action (4.20) appears by expressing the NR gauge fields in terms of the relativistic ones through the semigroup elements as

$$\begin{aligned}
 \omega &= \lambda_0 W^0, \quad \omega^a = \lambda_1 W^a, \quad s = \lambda_2 W^0, \\
 h &= \lambda_0 E^0, \quad e^a = \lambda_1 E^a, \quad m = \lambda_2 E^0, \\
 \psi^+ &= \lambda_0 \Psi^+, \quad \psi^- = \lambda_1 \Psi^-, \quad \rho = \lambda_2 \Psi^+, \\
 t_1 &= \lambda_0 \mathcal{A}, \quad t_2 = \lambda_2 \mathcal{A}, \\
 u_1 &= \lambda_0 \mathcal{C}, \quad u_2 = \lambda_2 \mathcal{C},
 \end{aligned}
 \tag{4.24}$$

plus the expansion of the relativistic parameters as in (4.15). Here, we have considered the decomposition of the Aindex

as in (3.4) along with the following redefinition of the spinor one-form fields:

$$\Psi^\pm = \frac{1}{\sqrt{2}} \left(\Psi_\alpha^1 \pm (\gamma^0)_{\alpha\beta} \Psi_\beta^2 \right). \quad (4.25)$$

5 Conclusions

In this work we have presented the NR counterpart of the teleparallel CS (super)gravity theory in three spacetime dimensions. We first focused on the purely bosonic case, studying the NR limit of the relativistic teleparallel CS gravity [17], based on the teleparallel algebra (2.1). A torsional NR algebra has been obtained by considering a particular $U(1)$ -enlargement of (2.1). Such enlargement, with involves additional Abelian generators, is required firstly to ensure a non-degenerate invariant tensor in the limit process to the NR theory and secondly to have a well-defined flat limit leading to the extended Bargmann symmetry. The NR action obtained is characterized, on-shell, by non-vanishing spatial torsion, which is sourced by the cosmological constant. In the vanishing cosmological constant limit $\ell \rightarrow \infty$ such spatial torsion vanishes and the NR CS gravity model boils down to the extended Bargmann gravity theory [54]. The latter was also obtained as the vanishing cosmological constant limit of the CS gravity theory invariant under the Newton–Hooke symmetry. However, as we have shown, the inclusion of the cosmological constant through the torsional NR algebra is different from the one constructed using the Newton–Hooke symmetry. Hence, we have also provided another NR CS gravity model which can be seen as an alternative cosmological extension of the extended Bargmann gravity theory.

At the supersymmetric level, we have considered the three-dimensional relativistic $\mathcal{N} = 2$ teleparallel CS supergravity theory [17] and developed its NR counterpart by performing the S-expansion procedure [60]. Indeed, the S-expansion method can be seen as a generalization of the Inönö–Wigner contraction when the semigroup considered belongs to the $S_E^{(N)}$ family. In particular, the semigroup $S_E^{(2)}$ allows to obtain not only a well-defined NR teleparallel (super)algebra but also its invariant tensor. Analogously to the non-vanishing super-torsion appearing in the relativistic three-dimensional $\mathcal{N} = 2$ teleparallel CS supergravity, the torsional NR CS supergravity theory presented here is characterized, on-shell, by a non-vanishing spatial super-torsion sourced by the cosmological constant. In the vanishing cosmological constant limit $\ell \rightarrow \infty$ of the aforementioned NR model, the supersymmetric extension of the extended Bargmann gravity theory [54, 57, 62] is recovered.

To obtain the respective NR version of a relativistic superalgebra is a challenging task and, to our knowledge, has only been approached in three spacetime dimensions [52–57, 62–66, 105]. The S-expansion method not only offers us a straightforward mechanism to derive its corresponding NR counterpart but also provides us with the non-vanishing components of the NR invariant tensor. Such procedure could be useful to go beyond three spacetime dimensions and to derive NR CS (super)gravity model in higher odd spacetime dimensions. In even dimensions, a completely different approach should be considered. In this direction, a geometrical formulation of supergravity à la MacDowell–Mansouri [106–108] could be used.

Further generalizations of the teleparallel CS gravity theory has been obtained in [12] by considering a non-vanishing torsion in the Maxwell CS gravity theory [109–117]. The study of black hole solutions, and their thermodynamics, of the Maxwellian teleparallel gravity could bring valuable information about the physical implications of a non-vanishing torsion in Maxwell gravity theory [work in progress]. On the other hand, a Maxwellian version of the teleparallel supergravity remains unexplored and could be useful not only to elucidate the role of torsion in Maxwell supergravity but also to approach a Maxwellian generalization of the present torsional NR supergravity. It would be then interesting to analyze the differences with our torsional NR (super)gravity theory. In particular, one could expect to find a deformation of the Maxwellian extended Bargmann (super)algebra [56, 102].

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